

Sc and neutron-capture abundances in Galactic low- and high- α field halo stars

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ABSTRACT

We determine relative abundance ratios for the neutron-capture elements Zr, La, Ce, Nd and Eu for a sample of 27 Galactic dwarf stars with $-1.5 < [\text{Fe}/\text{H}] < -0.8$. We also measure the iron-peak element Sc. These stars separate into three populations (low- and high- α halo and thick-disc stars) based on the $[\alpha/\text{Fe}]$ abundance ratio and their kinematics as discovered by Nissen & Schuster. We find differences between the low- and high- α groups in the abundance ratios of $[\text{Sc}/\text{Fe}]$, $[\text{Zr}/\text{Fe}]$, $[\text{La}/\text{Zr}]$, $[\text{Y}/\text{Eu}]$ and $[\text{Ba}/\text{Eu}]$ when including Y and Ba from Nissen & Schuster. For all ratios except $[\text{La}/\text{Zr}]$, the low- α stars have a lower abundance compared to the high- α stars. The low- α stars display the same abundance patterns of high $[\text{Ba}/\text{Y}]$ and low $[\text{Y}/\text{Eu}]$ as observed in present-day dwarf spheroidal galaxies, although with smaller abundance differences, when compared to the high- α stars. These distinct chemical patterns have been attributed to differences in the star formation rate between the two populations and the contribution of low-metallicity, low-mass asymptotic giant branch (AGB) stars to the low- α population. By comparing the low- α population with AGB stellar models, we place constraints on the mass range of the AGB stars.

Key words: stars: abundances – Galaxy: halo.

1 INTRODUCTION

Understanding the earliest phases of the assembly of our Galaxy, the Milky Way remains a key issue in contemporary astronomy. Cosmological simulations predict that mergers and accretion of satellite dwarf galaxies may play an important role through hierarchical structure formation (e.g. Bullock & Johnston 2005; Pillepich et al. 2014). Our Galaxy is no exception. Observational support for the hierarchical nature of Galactic formation, and the role dwarf galaxies play in that process, comes from the accretion of the Sagittarius dwarf galaxy (Ibata, Gilmore & Irwin 1994) and numerous tidal streams observed in the Galactic halo (e.g. Belokurov et al. 2006).

The chemical abundances and kinematics of Galactic halo stars hold valuable clues concerning Galactic formation at the earliest times (e.g. Eggen, Lynden-Bell & Sandage 1962; Searle &

Zinn 1978; Venn et al. 2004; Frebel & Norris 2015). Due to their long lifetimes, FGK-type dwarf stars are ideal for tracing the evolution of the Galaxy. The atmospheres of dwarf stars closely mirror the chemical composition of the gas clouds from which they formed. This means that spectroscopic studies of these ‘stellar fossils’ will enable us to probe the chemical and kinematic history of our Galaxy over cosmic time.

Comparison of chemical abundances between the Galactic halo and satellite dwarf spheroidal galaxies (dSph) reveals differences in their chemical composition, most notably in the $[\alpha/\text{Fe}]$ ratio (e.g. Shetrone, Côté & Sargent 2001; Tolstoy et al. 2003; Venn et al. 2004; Tolstoy, Hill & Tosi 2009). On average, dSph stars with $[\text{Fe}/\text{H}] \gtrsim -2$ have lower $[\alpha/\text{Fe}]$ than the Galaxy. The $[\alpha/\text{Fe}]$ ratio depends on the relative number of Type Ia to Type II supernovae, and this quantity is sensitive to the star formation rate (e.g. Tinsley 1979; Matteucci & Greggio 1986; Wheeler, Sneden & Truran 1989; Nomoto, Kobayashi & Tominaga 2013). The lower $[\alpha/\text{Fe}]$ observed in dSph would therefore suggest a slower star formation rate relative to the Galaxy. Differences in the $[\alpha/\text{Fe}]$ ratio

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between the halo and present-day dwarf galaxies preclude the continuous merging of such systems, although it may be that the mergers of massive dwarf galaxies occurred at very early times in the formation of the Galactic halo (Venn et al. 2004).

Although a handful of α -poor stars were known to exist in the Galactic halo (e.g. Carney et al. 1997; King 1997; Preston & Snedden 2000; Ivans et al. 2003), the study by Nissen & Schuster (2010) (hereafter NS10) provided the first compelling evidence that the halo hosts two populations that are well separated in $[\alpha/\text{Fe}]$. NS10 studied 94 halo dwarf stars in the solar neighbourhood using a careful differential analysis resulting in abundance errors for $[\alpha/\text{Fe}]$ of only 0.02 dex. They argued that the low- α stars have kinematics consistent with accretion from a satellite galaxy with a slower star formation rate than the high- α population that most likely formed *in situ* in the Galaxy. Additionally, NS10 found that the low- and high- α populations separate in the $[\text{Na}/\text{Fe}]$ versus $[\text{Ni}/\text{Fe}]$ abundance plane. No separation was seen for $[\text{Cr}/\text{Fe}]$.

The study of NS10 was extended by Ramírez, Meléndez & Chanamé (2012). Through an analysis of the 777 nm O I triplet lines, they found that the low- α stars have a lower $[\text{O}/\text{Fe}]$ abundance compared to the high- α and thick-disc (TD) stars. The low- α stars also show a cosmic star-to-star scatter in $[\text{O}/\text{Fe}]$ at a given $[\text{Fe}/\text{H}]$.

Before continuing, it is important to recognize that any attempts to infer the chemical enrichment history of a stellar population requires a detailed understanding of the nucleosynthetic origin of individual elements (Nissen 2013). Low- to intermediate-mass stars undergo rich nucleosynthesis once they reach the asymptotic giant branch (AGB) phase. It has been observationally confirmed that the slow neutron-capture process (*s*-process) takes place in AGB stars (e.g. Wallerstein et al. 1997; Busso, Gallino & Wasserburg 1999). The *s*-process is responsible for the production of around half of the abundance of the heavy elements beyond iron. Through strong mass-loss, AGB stars eject enriched material into the interstellar medium and pollute the next generation of stars (Herwig 2005; Karakas & Lattanzio 2014).

AGB stars may play an important role in the origin of chemical abundance differences between the Galaxy and nearby dwarf galaxies. In addition to $[\alpha/\text{Fe}]$, differences in the ratios of $[\text{Ba}/\text{Y}]$ and $[\text{Y}/\text{Eu}]$ are found between dSph and the Galaxy (e.g. Tolstoy et al. 2009). The high $[\text{Ba}/\text{Y}]$ observed in dSph was attributed to low-mass AGB stars enriching the interstellar medium as a result of the slower chemical evolution. To explain the low $[\text{Y}/\text{Eu}]$ observed, Venn et al. (2004) suggested that metal-poor AGB stars did not contribute as much Y (compared to more metal-rich AGB stars) because the production of second *s*-process peak elements such as Ba and La were favoured over first *s*-process peak elements such as Zr and Y (Busso et al. 2001).

Nissen & Schuster (2011) (hereafter NS11) determined abundances for two neutron-capture elements produced by the *s*-process in AGB stars, Ba and Y (along with the iron-peak elements Mn, Cu and Zn). They discovered that the low- and high- α populations could be separated in $[\text{Cu}/\text{Fe}]$ and $[\text{Zn}/\text{Fe}]$. More importantly, NS11 found that the $[\text{Ba}/\text{Y}]$ ratio differs between the two populations with the low- α stars showing a higher $[\text{Ba}/\text{Y}]$ value. Those chemical abundance patterns are similar to that found in dSph by Venn et al. (2004) and reinforces the hypothesis that the low- α stars are likely to have been accreted from dSph.

The NS10 sample provides a unique opportunity to study additional neutron-capture element abundances in the low- and high- α populations. Such an analysis will enable us to examine the abundances for a larger range of neutron-capture elements, with

high-precision, in stars likely to have been accreted from dwarf galaxies. Stars currently in dwarf galaxies are very faint such that high-precision chemical abundance studies cannot be conducted with existing facilities. The aim of this study is to investigate the role of AGB stars in the chemical enrichment of the Galactic halo and dSph. Specifically, we build upon the work of NS10 and NS11 by studying additional *s*-process elements (Zr, La, Ce and Nd). We also measure, for the first time with this unique sample, a rapid neutron-capture process (*r*-process) element (Eu). Finally, we measure Sc that has been observed to behave like an α element (e.g. Nissen et al. 2000).

2 SAMPLE SELECTION AND OBSERVATIONS

Our observations focused on obtaining high-quality blue spectra for stars in the NS10 sample. In the analyses of NS10 and NS11, only lines redder than 4700 Å were measured. Our targets were selected to include both low- and high- α stars across the metallicity range of $-1.5 < [\text{Fe}/\text{H}] < -0.8$. The sample consists of 13 low- α stars, 6 high- α stars and 8 TD stars, including the two reference TD stars used in the studies by NS10 and NS11, HD 22879 and HD 76932.

Programme stars were observed using the Magellan Inamori Kyocera Echelle (MIKE) spectrograph (Bernstein et al. 2003) on the Magellan (Clay) telescope during 2011 February and June. We used the 0.35 arcsec slit size that provides a resolving power of 83 000 in the blue. Typical exposure times ranged from 500 to 5000 s per target and the signal-to-noise ratio per pixel of the spectra was approximately 400 at 4500 Å. The wavelength coverage provided by MIKE is 3350–9000 Å. Data reduction was accomplished using IRAF¹ with the MIKE MTOOLS² package. Continuum fitting was performed using the Spectroscopy Made Hard (SMH) package as used in Casey et al. (2014) and described in Casey (2014).

3 ABUNDANCES

Following NS10, we perform a differential line-by-line analysis relative to the TD star HD 22879. This approach enables us to obtain high-precision abundance ratios that are necessary in identifying possible separations between the low- and high- α populations. The model atmospheres used in the analysis were one-dimensional, plane parallel, local thermodynamic equilibrium (LTE) ATLAS9 models by Castelli & Kurucz (2003). Interpolation between T_{eff} , $\log g$, $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$ was required to produce model atmospheres. The stellar parameters, T_{eff} , ξ_{turb} and $\log g$, as well as the abundance ratios of $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$, were taken from NS10 and NS11 (see Table 1). For a complete discussion on the determination of these parameters, see NS10 and NS11. We use stellar parameters from NS10 without the correction for effective temperature of ~ 100 K from NS11. Given the differential nature of the abundance analysis and that most stars have fairly similar stellar parameters, a systematic shift of 100 K is likely to have a minimal effect on the final differential abundances.

Elemental abundances for Zr, Ce and Nd were derived from the analysis of equivalent widths (EWs) using SMH. For lines with

¹ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy under a cooperative agreement with the National Science Foundation.

² www.lco.cl/telescopes-information/magellan/instruments/mike/iraf-tools/iraf-mtools-package

Table 1. Atmospheric parameters and abundance ratios for the programme stars (taken from NS10 and NS11).

ID	T_{eff} (K)	$\log g$ (cgs)	ξ_{turb} (km s^{-1})	[Fe/H]	$[\alpha/\text{Fe}]$	Class ^a
CD -45 3283	5597	4.55	1.0	-0.91	0.12	low- α
CD -61 0282	5759	4.31	1.3	-1.23	0.22	low- α
G 16-20	5625	3.64	1.5	-1.42	0.26	(low- α)
G 18-28	5372	4.41	1.0	-0.83	0.31	high- α
G 18-39	6040	4.21	1.5	-1.39	0.34	high- α
G 31-55	5638	4.30	1.4	-1.10	0.29	high- α
G 56-30	5830	4.26	1.3	-0.89	0.11	low- α
G 56-36	5933	4.28	1.4	-0.94	0.20	low- α
G 66-22	5236	4.41	0.9	-0.86	0.12	low- α
G 170-56	5994	4.12	1.5	-0.92	0.17	low- α
HD 3567	6051	4.02	1.5	-1.16	0.21	low- α
HD 22879	5759	4.25	1.3	-0.85	0.31	TD
HD 25704	5868	4.26	1.4	-0.85	0.24	TD
HD 76932	5877	4.13	1.4	-0.87	0.29	TD
HD 103723	5938	4.19	1.2	-0.80	0.14	low- α
HD 105004	5754	4.30	1.2	-0.82	0.14	low- α
HD 111980	5778	3.96	1.5	-1.08	0.34	high- α
HD 120559	5412	4.50	1.1	-0.89	0.30	TD
HD 126681	5507	4.45	1.2	-1.17	0.35	TD
HD 179626	5850	4.13	1.6	-1.00	0.32	high- α
HD 189558	5617	3.80	1.4	-1.12	0.35	TD
HD 193901	5650	4.36	1.2	-1.07	0.17	low- α
HD 194598	5942	4.33	1.4	-1.09	0.20	low- α
HD 199289	5810	4.28	1.3	-1.04	0.30	TD
HD 205650	5698	4.32	1.3	-1.17	0.30	TD
HD 219617	5862	4.28	1.5	-1.45	0.28	(low- α)
HD 230409	5318	4.54	1.1	-0.85	0.27	high- α

Note. ^aFor stars with $[\text{Fe}/\text{H}] < -1.4$, the classification is uncertain (NS10).

Table 2. EW measurements (in mÅ) of Zr II, Ce II and Nd II for each star.

ID	Zr II 4208.98 Å	Ce II 4562.36 Å	Ce II 4628.16 Å	Nd II 4706.54 Å
CD -45 3283	21.3	7.6	5.3	7.0
CD -61 0282	17.7	5.2	2.7	4.0
G 16-20	22.1	6.6	4.1	4.0
G 18-28	29.1	10.3	...	9.4
G 18-39	17.3	...	2.0	3.0
G 31-55	23.8	4.9	3.6	4.6
G 56-30	23.1	6.6	7.2	2.9
G 56-36	26.1	6.7	4.4	5.4
G 66-22	26.2	...	8.9	10.6
G 170-56	23.4	6.8	3.4	4.2
HD 3567	21.3	4.9	5.3	4.0
HD 22879	31.4	8.2	6.5	7.9
HD 25704	22.4	6.5	6.1	7.1
HD 76932	34.3	8.4	5.1	7.9
HD 103723	27.5	9.4	7.4	7.7
HD 105004	24.6	...	9.4	5.0
HD 111980	37.4	10.6	6.5	6.6
HD 120559	23.8	...	4.1	7.5
HD 126681	28.9	7.4	5.4	5.0
HD 179626	27.6	5.9	4.1	4.1
HD 189558	39.3	10.9	8.9	7.1
HD 193901	17.4	5.2	...	5.3
HD 194598	15.8	3.6	3.2	2.9
HD 199289	19.1	3.3	2.5	3.6
HD 205650	19.1	4.6	3.4	3.0
HD 219617	10.7	...	3.3	2.0 ^a
HD 230409	26.2	8.2	5.8	10.5

Note. ^aUpper limit.

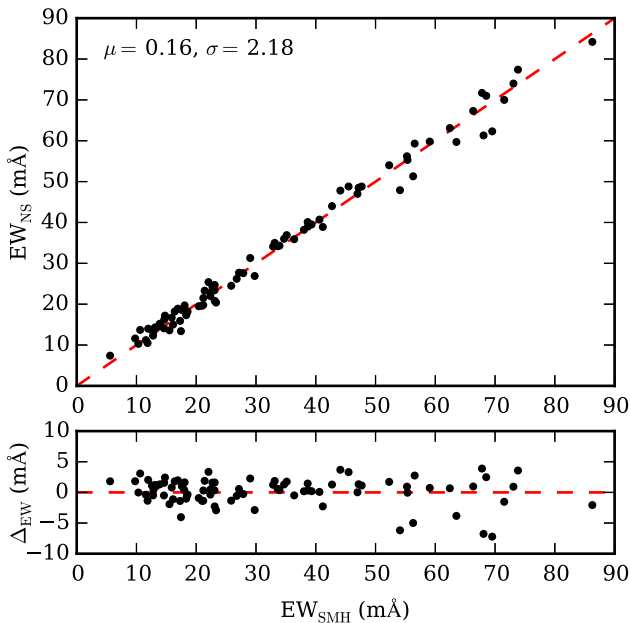


Figure 1. Comparison between EW measurements for HD 22879 from NS10 and NS11 (EW_{NS}) and by using SMH (EW_{SMH}, this work). The dashed line shows the one-to-one correspondence line. The bottom panel shows the difference between the two analyses.

EW < 60 mÅ, a Gaussian profile was fitted to the line. For lines with EW larger than 60 mÅ, a Voigt profile was used. Fig. 1 presents a comparison of 85 EW measurements (excluding Fe I and Fe II) of NS10 and NS11 to the EWs measured using SMH with the same

line list for HD 22879. The mean EW difference between the two methods is 0.16 ± 0.24 mÅ ($\sigma = 2.18$ mÅ).

We use the atomic data compiled by Yong et al. (2005) for the measured absorption lines. The abundance of Zr is derived from the EW of the Zr II line at 4208.98 Å, which is unblended with other lines, allowing for a reliable measurement. The Zr II line is detected in all stars with all EW measurements being greater than 10 mÅ (see Table 2). Ce abundances have been determined using the EWs of two Ce II lines at 4562.36 and 4628.16 Å. Seven stars only have one Ce II line available for measurement: five stars do not have a measurable 4562.36 Å line whereas two stars do not have a measurable 4628.16 Å line (see Table 2). For all but three Ce lines, the EW is less than 10 mÅ.

Nd abundances have been determined using the Nd II line at 4706.54 Å. The EW measurements are less than 10 mÅ except in two stars (G 66-22 and HD 230409). We arbitrarily set a minimum reliable EW measurement limit of 2 mÅ. One star, the most metal-poor in the sample (HD 219617), falls below our 2 mÅ limit and therefore we adopt an upper limit of 2 mÅ for the Nd EW.

For Sc, La and Eu, we fit synthetic spectra generated by the LTE spectrum synthesis program MOOG (v. 2013, Sneden 1973) to determine abundances. These elements require the consideration of hyperfine splitting and/or isotopic shifts. Fig. 2 illustrates the fit of the synthetic spectrum to the observed spectrum for the Sc and Eu lines of HD 22879. The best-fitting synthetic spectrum was determined using χ^2 minimization. For Sc, the $\log gf$ values of Nissen et al. (2000) were utilized for the 5526.81 Å absorption line. La was determined using one La II line at 4086.71 Å with $\log gf$ values from Lawler, Bonvallet & Sneden (2001). For the Eu II spectral line at 4129.72 Å, we use the $\log gf$ values from Lawler et al. (2001) and solar isotope ratios from Lodders (2003).

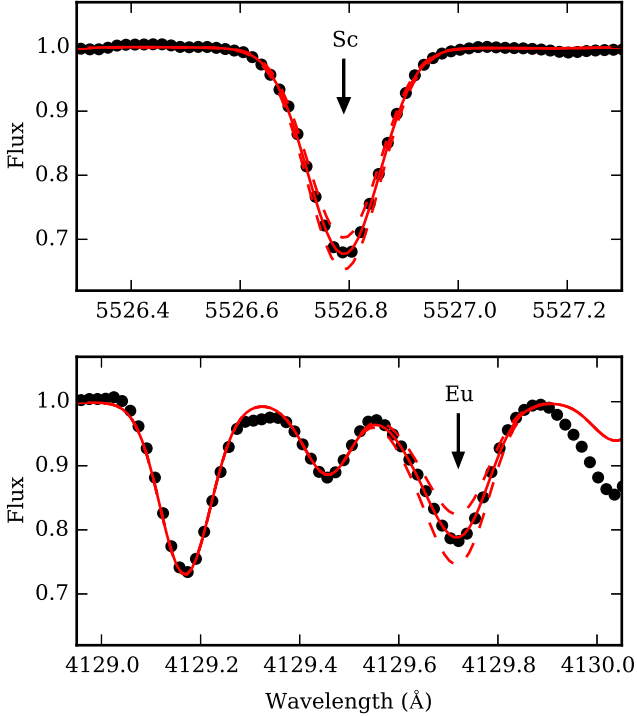


Figure 2. Synthetic spectrum fit (solid line) to the observations (black points) of the Sc line at 5526 Å (top) and the Eu line at 4129 Å (bottom) for HD 22879. The dashed lines show synthetic spectra with $[\text{Sc}/\text{Fe}]$ and $[\text{Eu}/\text{Fe}]$ of ± 0.1 .

Table 3. Adopted solar elemental abundances from Asplund et al. (2009).

Element	Z	$\log(X/\text{H}) + 12$
Sc ^a	21	3.05
Zr	40	2.58
La	57	1.10
Ce	58	1.58
Nd	60	1.42
Eu	63	0.52

Note. ^aMeteoritic value.

We adopt solar elemental abundances from Asplund et al. (2009) as detailed in Table 3. Since we perform a differential analysis to measure abundances, the choice of solar abundance only shifts the stellar abundances up or down and does not affect the relative abundances. To reiterate, we are primarily interested in relative abundances, and therefore abundance differences, between the low- and high- α populations. The abundances are given in Table 4.

We estimated errors using the stellar parameter uncertainties from NS10 where the differential errors for $\sigma(\log g)$ and $\sigma(T_{\text{eff}})$ are ± 0.05 and ± 30 K, respectively. The 1σ errors from NS10 for $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$ were 0.04 and 0.02 dex, respectively. We estimated the 1σ errors in $[\text{Zr}/\text{Fe}]$ and $[\text{Ce}/\text{Fe}]$ to be 0.04. For $[\text{Sc}/\text{Fe}]$, $[\text{La}/\text{Fe}]$ and $[\text{Eu}/\text{Fe}]$, the errors are 0.05. The error for $[\text{Nd}/\text{Fe}]$ is 0.07.

3.1 Zr

The neutron-capture element Zr is predominately produced by AGB stars and is a first s -process peak element. The s -process contribution to Zr in the Solar system is ~ 80 per cent (e.g. Simmerer et al. 2004; Sneden, Cowan & Gallino 2008; Bisterzo et al. 2014).

By determining the abundance of Zr, we expand on the measurements of Y, another first s -process peak element whose abundances were previously obtained by NS11.

Fig. 3(a) shows $[\text{Zr}/\text{Fe}]$ as a function of $[\text{Fe}/\text{H}]$. The dispersion in $[\text{Zr}/\text{Fe}]$ is ~ 0.12 dex, which is larger than the error of 0.04 dex. When considering the two populations, the scatter in $[\text{Zr}/\text{Fe}]$ is smaller for the low- α stars with a dispersion of 0.06 dex compared to a dispersion of 0.13 dex when combining the high- α and TD stars. In addition to the difference in the dispersion of $[\text{Zr}/\text{Fe}]$, the low- α stars have a systematically lower $[\text{Zr}/\text{Fe}]$ abundance compared to the high- α and TD stars.

This possible separation in the high- and low- α populations is in agreement with the slight separation in abundance measurements of $[\text{Y}/\text{Fe}]$ for the two populations seen in NS11. As Zr and Y are both first-peak s -process elements, they are expected to follow the same trend. Fig. 4 demonstrates the difference in the Zr abundance for a low- α star (HD 194598) and a high- α star (HD 111980) at nearly the same metallicity (see Table 1). The low- α star has a weaker Zr II line than the high- α star. While the effective temperature differs between these two stars by approximately 150 K, all other lines in this region exhibit very similar strengths.

A comparison between $[\text{Y}/\text{Fe}]$ as measured by NS11 and $[\text{Zr}/\text{Fe}]$ is presented in Fig. 5. The abundance of Zr follows the abundance of Y, which is expected if Y and Zr have a common origin. While there is a systematic offset between $[\text{Y}/\text{Fe}]$ and $[\text{Zr}/\text{Fe}]$ that may be due to errors in the gf values of the lines, the important point is that the slope of the data is close to 1. The high- α stars have a tighter correlation between $[\text{Y}/\text{Fe}]$ and $[\text{Zr}/\text{Fe}]$ compared to the low- α stars.

3.2 La, Ce and Nd

We extend upon the Ba measurements of NS11 by determining abundances for other elements belonging to the second s -process peak, namely La, Ce and Nd.

The s -process contribution to La in the Solar system is 75 per cent and there are a number of advantages with using La as an s -process indicator compared to Ba (e.g. Simmerer et al. 2004; Winckler et al. 2006). While Ba consists of five stable isotopes and requires the consideration of isotope shifts, 99.91 per cent of La is in the isotope ^{139}La (Lodders 2003; Asplund et al. 2009). Fig. 3(b) presents $[\text{La}/\text{Fe}]$ as a function of $[\text{Fe}/\text{H}]$. There is no clear difference in the La abundance between the low- α , high- α and TD stars. The dispersion in the abundance of $[\text{La}/\text{Fe}]$ is ~ 0.07 dex, the lowest of all the s -process abundances measured here. The $[\text{La}/\text{Fe}]$ ratio is almost constant with $[\text{Fe}/\text{H}]$ and the mean abundance close to the solar value.

The s -process contribution to Ce in the Solar system is 81 per cent where ^{140}Ce is the only stable isotope accessible by the s -process (Sneden et al. 2008). Fig. 3(c) presents the results for $[\text{Ce}/\text{Fe}]$. The majority of stars have a subsolar $[\text{Ce}/\text{Fe}]$ abundance with a mean value of around -0.04 dex and no apparent separation between the low- and high- α stars. The dispersions for the low- α and high- α stars are essentially the same (0.09 dex compared to 0.10 dex, respectively).

For Nd, 58 per cent of the Solar system abundance is due to the s -process (Sneden et al. 2008). Fig. 3(d) presents the abundance of $[\text{Nd}/\text{Fe}]$ against $[\text{Fe}/\text{H}]$. Due to the weak Nd II line for HD 219617, we provide an upper limit of $[\text{Nd}/\text{Fe}] = 0.03$ for this object based on an EW limit of $2 \text{ m}\text{\AA}$. There is no obvious difference in $[\text{Nd}/\text{Fe}]$ between the two populations except for a few low- α stars having a lower Nd abundance. This causes the dispersion to be higher for

Table 4. Abundances for each of the measured elements. The values for [Fe/H] are taken from NS10.

Star	[Fe/H]	[Sc/Fe]	[Zr/Fe]	[La/Fe]	[Ce/Fe]	[Nd/Fe]	[Eu/Fe]
CD -45 3283	-0.91	-0.11	-0.07	0.00	-0.01	0.36	0.31
CD -61 0282	-1.23	-0.15	0.00	-0.01	-0.05	0.32	0.26
G 16-20	-1.42	-0.11	-0.00	0.02	-0.04	0.17	0.32
G 18-28	-0.83	0.11	0.03	-0.02	0.04	0.34	0.13
G 18-39	-1.39	-0.00	0.21	0.07	-0.03	0.47	0.26
G 31-55	-1.10	-0.03	0.06	-0.08	-0.11	0.26	0.05
G 56-30	-0.89	-0.22	-0.12	-0.06	-0.04	-0.09	0.09
G 56-36	-0.94	-0.02	0.04	-0.04	-0.05	0.31	0.13
G 66-22	-0.86	-0.16	-0.08	-0.04	0.02	0.34	0.30
G 170-56	-0.92	-0.14	-0.10	-0.02	-0.16	0.14	0.16
HD 3567	-1.16	-0.04	0.05	0.16	0.07	0.33	0.39
HD 22879	-0.85	0.12	0.06	-0.09	-0.04	0.34	0.16
HD 25704	-0.85	0.09	-0.16	-0.02	-0.09	0.33	0.15
HD 76932	-0.87	0.15	0.13	0.03	-0.08	0.36	0.21
HD 103723	-0.80	-0.03	-0.05	-0.03	0.01	0.33	0.17
HD 105004	-0.82	-0.04	-0.13	-0.12	0.10	0.10	0.04
HD 111980	-1.08	0.00	0.25	0.13	0.09	0.33	0.14
HD 120559	-0.89	0.17	-0.05	-0.05	-0.22	0.32	0.20
HD 126681	-1.17	0.00	0.28	0.14	0.13	0.34	0.20
HD 179626	-1.00	0.12	0.03	0.02	-0.14	0.14	0.03
HD 189558	-1.12	0.09	0.24	0.07	0.08	0.24	0.17
HD 193901	-1.07	-0.19	-0.14	-0.00	-0.08	0.31	0.24
HD 194598	-1.09	-0.07	-0.11	-0.03	-0.11	0.17	0.27
HD 199289	-1.04	0.08	-0.09	-0.12	-0.30	0.15	0.09
HD 205650	-1.17	-0.01	0.01	-0.03	-0.07	0.14	0.19
HD 219617	-1.45	-0.11	-0.05	-0.11	0.18	0.03 ^a	0.21
HD 230409	-0.85	0.05	-0.03	-0.00	-0.07	0.43	0.27

Note. ^aUpper limit.

the low- α stars relative to the high- α stars, i.e. 0.14 dex compared to 0.10 dex. The mean abundance of [Nd/Fe] is 0.26 dex and all but one star have a supersolar [Nd/Fe] abundance.

3.3 Eu

Eu is predominately an r -process element with less than 2 per cent of the Solar system abundance a result of s -process nucleosynthesis (Snedden et al. 2008). We have included Eu in this study to investigate the contribution of r -process nucleosynthesis to the low- and high- α stars.

Fig. 6(a) presents the [Eu/Fe] abundances where there is a potential separation between the two populations: the low- α stars appear to show a higher [Eu/Fe] abundance compared to the high- α stars. However, due to the lack of stars at lower metallicities, the separation is not as clear as what is found for other elements such as Zr. The low- α group also shows a potential decrease in [Eu/Fe] with increasing metallicity. A larger sample of stars is needed to confirm these results.

3.4 Sc

There is only one stable isotope of Sc (^{45}Sc) and it is mainly synthesized during carbon and neon burning in massive stars (Woosley, Heger & Weaver 2002). It is also produced to a lesser extent in AGB stars (e.g. Fishlock et al. 2014). We measure Sc in order to verify the conclusions of Nissen et al. (2000) in which the low- α stars have a lower [Sc/Fe] abundance compared to the high- α stars.

Fig. 6(b) reveals a clear separation in [Sc/Fe]. The low- α stars have a subsolar [Sc/Fe] abundance with a mean value of -0.11 dex, whereas the high- α stars have a mostly supersolar abundance with

a mean value of 0.07 dex. These results confirm the separation in [Sc/Fe] found in the earlier work by Nissen et al. (2000) between the low- and high- α populations.

4 ANALYSIS

As with the NS10 and NS11 studies, we combine the TD and high- α halo stars into one high- α group. Due to a smaller sample size compared to NS11, we do not remove the abundance dependence on [Fe/H] for the high- α stars, as it cannot be definitively determined. Given that our sample size is smaller than in NS10, we now seek to quantify this potential limitation by introducing our methodology on how we decide whether a given abundance ratio exhibits a difference (i.e. separation) between the low- and high- α groups.

We first consider the NS10 measurements of Cr, a representative element for which NS10 found no significant separation between the low- and high- α groups. Using a two sample Kolmogorov–Smirnov (KS) test between the low- and high- α groups (without distinguishing TD stars), we find a p -value of 0.07 when considering [Cr/Fe] for all stars in the NS10 sample with $-1.45 \leq [\text{Fe}/\text{H}] \leq -0.8$.

In order to investigate how our smaller sample size affects our ability to identify abundance separations, we performed the following tests. We randomly select 13 low- α and 14 high- α stars from the full NS10 sample with $-1.45 \leq [\text{Fe}/\text{H}] \leq -0.8$, apply the KS test and obtain the p -value. We repeat this exercise 10 000 times and find that for approximately 25 per cent of the realizations, the p -value is below 0.07. If we arbitrarily define the threshold to be a p -value of 0.03, roughly 10 per cent of the random realizations result in a p -value below 0.03.

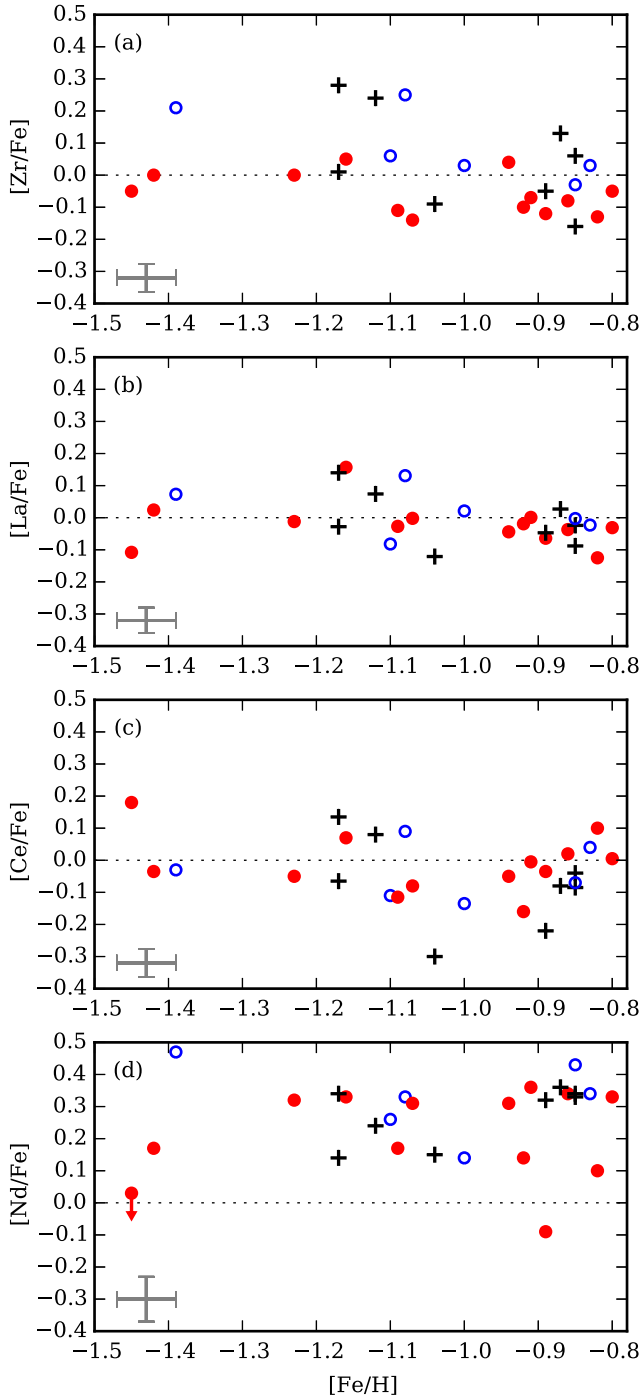


Figure 3. Four s -process elements (Zr, La, Ce and Nd) versus $[\text{Fe}/\text{H}]$. The red circles indicate low- α stars, the blue open circles indicate high- α stars and the plus signs indicate TD stars. The 1σ error is shown in the bottom left of each plot.

We now examine $[\text{Ba}/\text{Y}]$, an abundance ratio that was found to exhibit a significant separation between the low- and high- α populations. For the full NS10 sample with $-1.45 \leq [\text{Fe}/\text{H}] \leq -0.8$, we compute a p -value of 2.9×10^{-9} . We then seek to quantify the effect of having a smaller sample size. As above, we randomly select 13 low- α and 14 high- α stars from the full NS10 sample with $-1.45 \leq [\text{Fe}/\text{H}] \leq -0.8$, apply the KS test and obtain the p -value. We repeat the exercise 10 000 times. For a threshold p -value of 0.03,

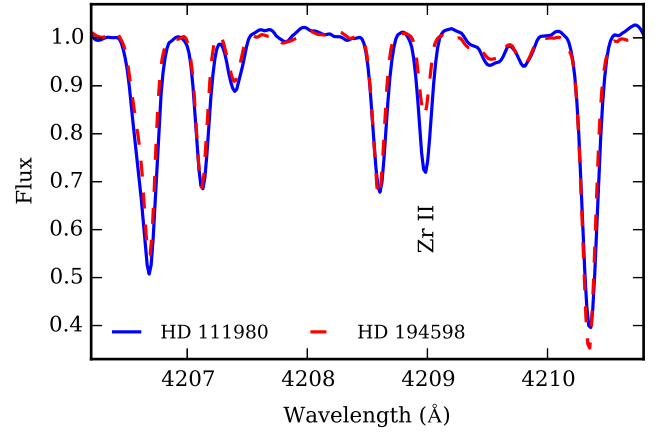


Figure 4. MIKE spectra around the Zr line at 4208.98 \AA for a low- α star (HD 194598) and a high- α star (HD 111980).

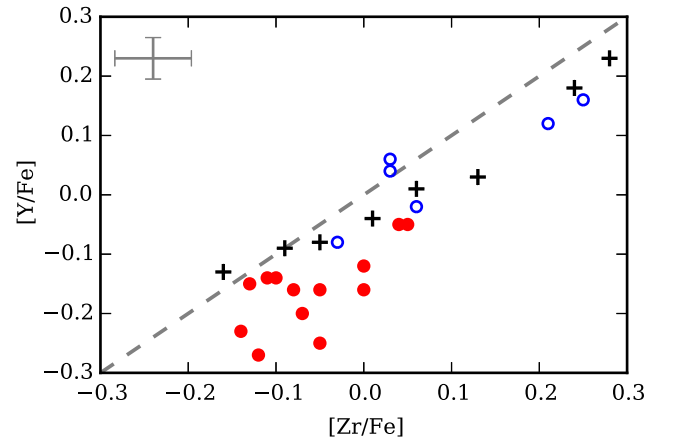


Figure 5. Comparison of $[\text{Y}/\text{Fe}]$ (taken from NS11) and $[\text{Zr}/\text{Fe}]$. The red circles indicate low- α stars, the blue open circles indicate high- α stars and the plus signs indicate TD stars. The dashed line shows the one-to-one correspondence line. The 1σ error is shown in the top left.

about 99.7 per cent of random realizations have a p -value below this level, indicating that we are able to identify the separation in $[\text{Ba}/\text{Y}]$ in all but 0.3 per cent of cases.

Based on these tests, we choose a p -value of 0.03, which ensures that there is at least a 90 per cent chance that the abundance separation in our sample of 27 stars is genuine despite having a smaller number of stars than NS10 and NS11. We also confirm using a KS test that the high- α population and TD stars have a significant probability of coming from the same population. In Table 5, we present the mean of each abundance ratio and its standard deviation for both low- and high- α groups. We also present the p -value of the two sample KS test between the low- and high- α groups. The analyses of $[\text{Y}/\text{Fe}]$, $[\text{Ba}/\text{Fe}]$ and $[\text{Ba}/\text{Y}]$ are included in Table 5 for comparison using the abundances measured by NS11 for our sample of 27 stars. For comparison, the mean p -value calculated for $[\text{Mg}/\text{Fe}]$ when randomly selecting 13 low- α and 14 high- α stars is 4.9×10^{-7} .

For abundances relative to Fe (e.g. $[\text{Zr}/\text{Fe}]$), Sc and Zr show a separation between the low- and high- α groups. For both elements, the high- α group has a higher $[\text{X}/\text{Fe}]$ ratio than the low- α group.

We find that the separation between the low- and high- α groups is detected for $[\text{Y}/\text{Fe}]$ despite having a smaller sample size than

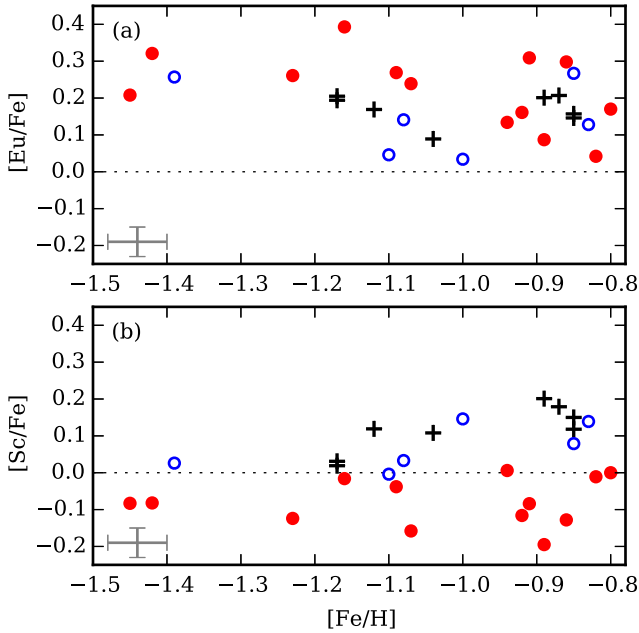


Figure 6. (a) $[\text{Sc}/\text{Fe}]$ and (b) $[\text{Eu}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$. The red circles indicate low- α stars, the blue open circles indicate high- α stars and the plus signs indicate TD stars. The 1σ error is shown in the bottom left of each plot.

Table 5. Mean value and standard deviation of each abundance ratio for the low- α group and the high- α group. Y and Ba are taken from NS11. The p -value between the two groups is also given.

	Low- α		High- α		p -value
	\bar{x}	σ	\bar{x}	σ	
$[\text{Sc}/\text{Fe}]$	-0.107	0.062	0.067	0.063	0.000
$[\text{Zr}/\text{Fe}]$	-0.058	0.061	0.069	0.131	0.017
$[\text{La}/\text{Fe}]$	-0.022	0.065	0.004	0.076	0.589
$[\text{Ce}/\text{Fe}]$	-0.012	0.087	-0.058	0.114	0.340
$[\text{Nd}/\text{Fe}]$	0.217	0.137	0.299	0.098	0.614
$[\text{Eu}/\text{Fe}]$	0.222	0.097	0.161	0.068	0.173
$[\text{La}/\text{Zr}]$	0.037	0.061	-0.066	0.084	0.005
$[\text{ls}/\text{Fe}]$	-0.109	0.057	0.049	0.117	0.004
$[\text{hs}/\text{Fe}]$	0.005	0.056	0.038	0.084	0.303
$[\text{hs}/\text{ls}]$	0.114	0.046	-0.010	0.065	0.000
$[\text{Y}/\text{Eu}]$	-0.382	0.101	-0.133	0.122	0.000
$[\text{ls}/\text{Eu}]$	-0.332	0.092	-0.112	0.126	0.001
$[\text{Ba}/\text{Eu}]$	-0.387	0.122	-0.253	0.122	0.024
$[\text{La}/\text{Eu}]$	-0.244	0.064	-0.157	0.082	0.173
$[\text{hs}/\text{Eu}]$	-0.218	0.082	-0.122	0.081	0.022
$[\text{Y}/\text{Fe}]$	-0.160	0.064	0.028	0.107	0.000
$[\text{Ba}/\text{Fe}]$	-0.165	0.070	-0.092	0.115	0.303
$[\text{Ba}/\text{Y}]$	-0.005	0.055	-0.120	0.030	0.000

NS11. In contrast to the first s -process peak elements Y and Zr, the abundance ratios for elements belonging to the second peak $[\text{La}/\text{Fe}]$, $[\text{Ce}/\text{Fe}]$ and $[\text{Nd}/\text{Fe}]$ show no significant difference between the low- and high- α groups. The second s -process peak element Ba, measured by NS11, also shows no separation in $[\text{Ba}/\text{Fe}]$ between the two groups.

We confirm the separation in $[\text{Ba}/\text{Y}]$ reported by NS11 (see Fig. 7a). In Fig. 7(b), we plot $[\text{La}/\text{Zr}]$ and are able to reproduce the separation between a light- s (ls) element at the first s -process peak and a heavy- s (hs) element at the second peak. The p -value of

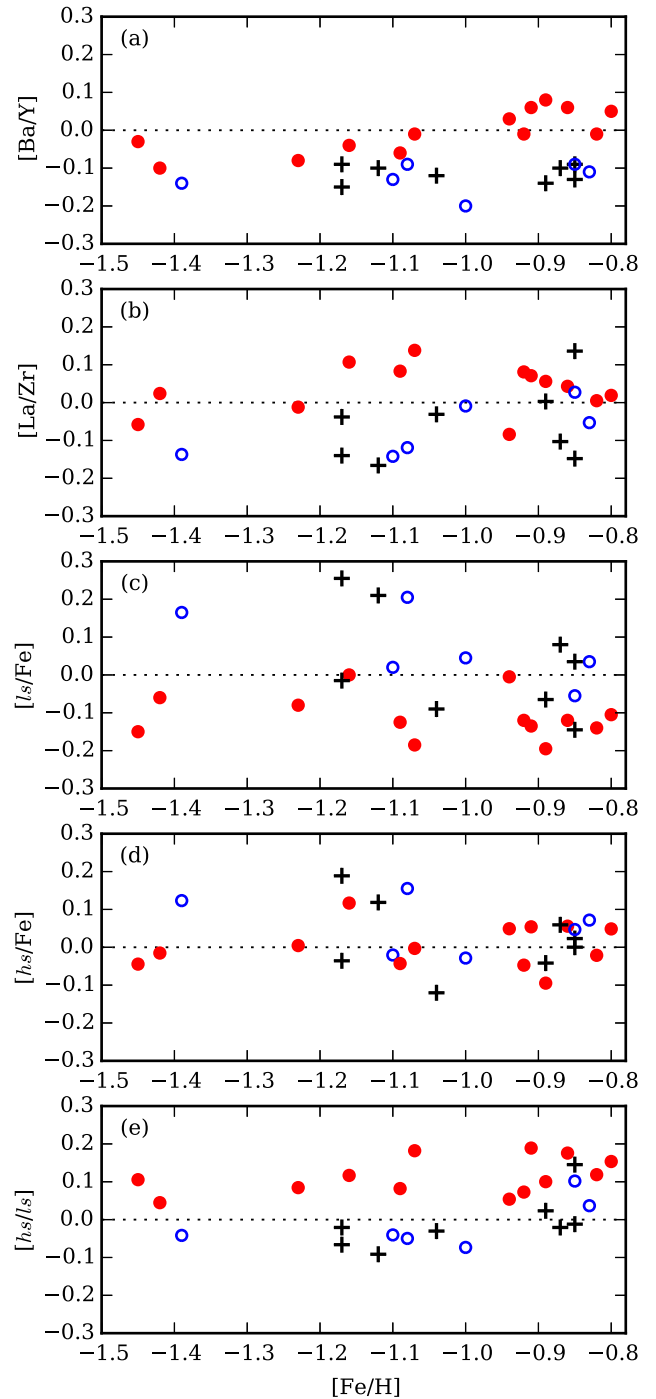


Figure 7. $[\text{Ba}/\text{Y}]$ (from NS11), $[\text{La}/\text{Zr}]$, $[\text{ls}/\text{Fe}]$, $[\text{hs}/\text{Fe}]$ and $[\text{hs}/\text{ls}]$ versus $[\text{Fe}/\text{H}]$, where ls is the average abundance of Y and Zr and hs is the average abundance of Ba, La, Nd and Ce. The red circles indicate low- α stars, the blue open circles indicate high- α stars and the plus signs indicate TD stars.

0.005 is lower than the individual p -values for $[\text{Zr}/\text{Fe}]$ (0.017) and $[\text{La}/\text{Fe}]$ (0.589) demonstrating that a clear separation between the low- and high- α populations can be seen when considering certain combinations of elements.

Assuming that the higher $[\text{La}/\text{Zr}]$ ratios in the low- α stars are due to enrichment from AGB stars, we use the model predictions of Fishlock et al. (2014) with $[\text{Fe}/\text{H}] = -1.2$ to estimate the mass range of those AGB stars. Starting at lower masses ($1\text{--}3 M_{\odot}$), calculations

predict $[La/Zr]$ ratios above solar; for example, the $3 M_{\odot}$ model predicts a final $[La/Zr]$ ratio of 0.32 dex. Moving to the intermediate masses (3.25 to $7 M_{\odot}$), calculations predict $[La/Zr]$ ratios below solar; for example, the $6 M_{\odot}$ model predicts a final $[La/Zr]$ ratio of -0.68 dex. Therefore, based on the $[La/Zr]$ ratio, AGBs stars in the range $1-3 M_{\odot}$ could be responsible for the abundance difference between the low- and high- α groups.

For stars in the range $-1.5 < [Fe/H] < -1$, there is a possible separation in $[Eu/Fe]$ where the low- α group has a higher $[Eu/Fe]$ abundance compared to the high- α group. However, higher metallicity stars show no difference between the two groups and the hint of a separation is not statistically significant with a p -value of 0.173. We return to this intriguing possibility of a separation in Section 5.

The s -process indicators are the average abundances of the elements at each of the two s -process peaks. We include the Y and Ba results from NS11 and define $[ls/Fe]$ and $[hs/Fe]^3$ to be:

$$[ls/Fe] = ([Y/Fe] + [Zr/Fe])/2 \quad (1)$$

and

$$[hs/Fe] = ([Ba/Fe] + [La/Fe] + [Ce/Fe] + [Nd/Fe])/4. \quad (2)$$

The $[ls/Fe]$ and $[hs/Fe]$ ratios against $[Fe/H]$ are presented in Figs 7(c) and (d). The low- α stars have a lower $[ls/Fe]$ ratio than the high- α stars. For $[hs/Fe]$, there is no apparent difference between the two populations.

Fig. 7(e) shows there is a clear separation in the s -process indicator $[hs/ls]$ between the low- and high- α groups. The low- α population has a higher $[hs/ls]$ ratio than the high- α population with a mean $[hs/ls]$ value of 0.11 dex compared to -0.01 dex for the high- α population. All the low- α stars have a supersolar $[hs/ls]$ ratio and this would be expected if the stars were enriched by previous generations of low-mass AGB stars. For example, the low-mass models of Fishlock et al. (2014) end the AGB phase with $[hs/ls]$ ratios of between 0.45 and 0.59 dex, excluding the $1 M_{\odot}$ model that experiences minimal third dredge-up. In contrast, the intermediate-mass models have final $[hs/ls]$ ratios between -0.71 and -0.41 dex. The differences between the low-mass and intermediate-mass models are a result of the different neutron sources (Lugaro et al. 2012).

The abundance ratio of s -process elements to r -process elements offers a powerful diagnostic for Galactic chemical evolution (Wheeler et al. 1989; McWilliam 1997). In Fig. 8, we present the abundance ratios of $[ls/Eu]$, $[Y/Eu]$, $[hs/Eu]$ and $[La/Eu]$ against $[Fe/H]$. The low- α stars have a lower $[ls/Eu]$ ratio compared to the high- α stars with the separation more evident at $[Fe/H] \lesssim -1$. There is also an apparent increase in $[ls/Eu]$ with increasing $[Fe/H]$ for the low- α group, which is not seen in the high- α group.

Fig. 8(b) presents the abundance ratio of the ls element, Y, relative to Eu. The low- α group has a lower abundance of $[Y/Eu]$ compared to the high- α group with a p -value of 2.2×10^{-4} . This p -value is more significant than the $[ls/Eu]$ separation. We also find that the $[Y/Eu]$ abundance for the low- α stars with $[Fe/H] < -1$ is almost constant at ~ -0.4 dex with minimal scatter. Similar behaviour is also seen for $[ls/Fe]$.

The separation in $[hs/Eu]$ is not as clear as that of $[ls/Eu]$, although it has a p -value of 0.022, just below the chosen significance level. The low- α group has a $[hs/Eu]$ ratio that is closer to the r -process

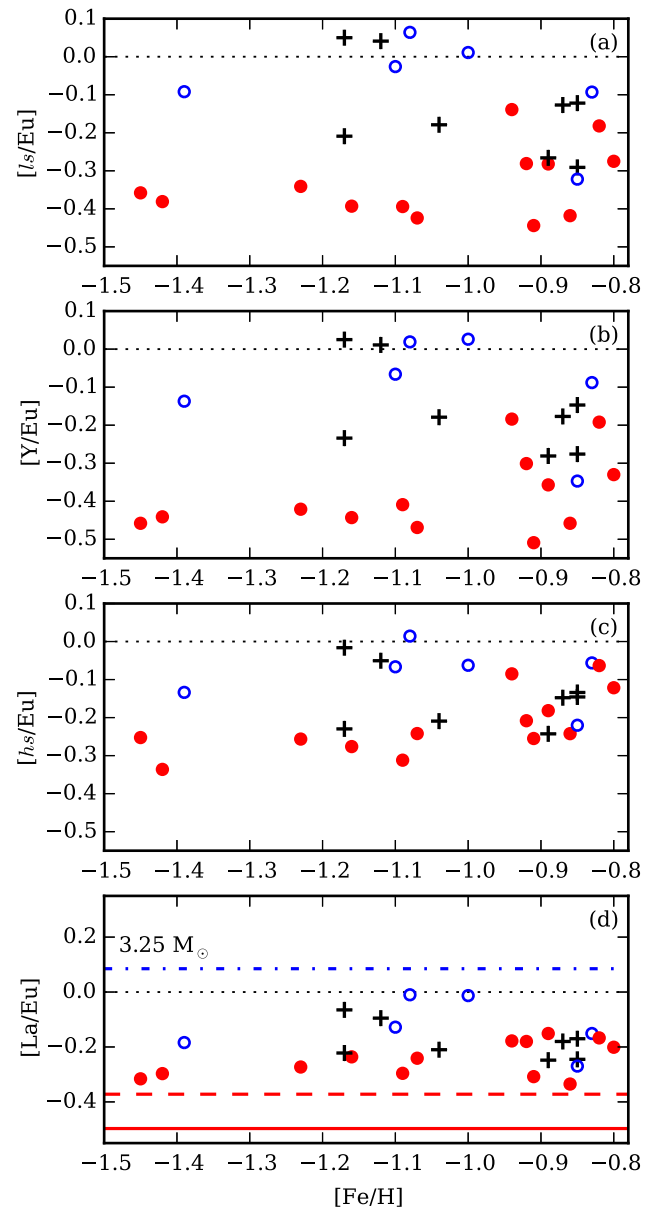


Figure 8. $[ls/Eu]$, $[Y/Eu]$, $[hs/Eu]$ and $[La/Eu]$ versus $[Fe/H]$, where ls is the average abundance of Y and Zr and hs is the average abundance of Ba, La, Nd and Ce. The red circles indicate low- α stars, the blue open circles indicate high- α stars and the plus signs indicate TD stars. The red solid line represents the r -process ratio for $[La/Eu]$ and the red dashed line its upper uncertainty (Winckler et al. 2006). The final surface abundance of $[La/Eu]$ for the $3.25 M_{\odot}$, $Z = 0.001$ AGB model from Fishlock et al. (2014) is also indicated (dash-dotted line).

contribution of the hs elements and Eu in the Solar system than the high- α group. As with $[ls/Fe]$, there is an apparent increase in $[hs/Eu]$ with increasing $[Fe/H]$ for the low- α stars.

Fig. 8(d) presents the abundance ratios of the hs element, La, relative to Eu where all the stars have a subsolar $[La/Eu]$ ratio. Unlike the separation found for $[hs/Fe]$, we do not find a separation in $[La/Eu]$ between the two groups. We also show the r -process contribution to the Solar system abundances of La and Eu (and its uncertainty) in Fig. 8(d) as determined by Winckler et al. (2006). The theoretical prediction of an AGB stellar model with an initial

³ As noted, the Solar system abundance of Nd has a ~ 40 per cent contribution from the r -process. Had we excluded Nd from hs , as did Lugaro et al. (2012), our conclusions would be unchanged.

mass of $3.25 M_{\odot}$ at a metallicity of $[\text{Fe}/\text{H}] = -1.2$, from Fishlock et al. (2014), is also presented in Fig. 8. Of all the models calculated by Fishlock et al. (2014), the lowest $[\text{La}/\text{Eu}]$ ratio of 0.08 dex occurs for the $3.25 M_{\odot}$ model.

5 DISCUSSION

We have studied chemical abundances in a subsample of low- and high- α stars from NS10 in which the low- α stars were likely accreted from satellite dwarf galaxies, whereas the high- α stars are representative of Galactic halo objects. We now seek to understand the abundance differences between these populations.

For Sc, we confirm and extend upon results from Nissen et al. (2000) in which the low- α stars have a lower $[\text{Sc}/\text{Fe}]$ ratio than the high- α stars. That is, Sc appears to behave like an α -capture element. The simplest explanation is that Sc is preferentially made in higher mass stars along with the α elements and that Type Ia SNe have a negligible contribution to Sc (Iwamoto et al. 1999; Woosley & Weaver 1995; Romano et al. 2010). Indeed, Nissen (2016) find that the $[\text{Sc}/\text{Fe}]$ ratio exhibits a similar behaviour with age as $[\text{Mg}/\text{Fe}]$ in their sample of solar twins. The low- α stars have low $[\text{Sc}/\text{Fe}]$ ratios due to the extra contribution of Fe from Type Ia SNe.

As noted in the Introduction, there is another chemical signature besides $[\alpha/\text{Fe}]^4$ that distinguishes halo stars from their dSph counterparts, namely, the neutron-capture element abundance ratios. In dSph, roughly two-thirds of the stars show a higher $[\text{Ba}/\text{Y}]$ ratio and a lower $[\text{Y}/\text{Eu}]$ ratio compared to Galactic halo stars (Venn et al. 2004; Tolstoy et al. 2009). McWilliam, Wallerstein & Mottini (2013) analysed three Sagittarius dSph stars and found low $[\alpha/\text{Fe}]$, $[\text{Na}/\text{Fe}]$, $[\text{Al}/\text{Fe}]$ and $[\text{Cu}/\text{Fe}]$ ratios along with a high $[\text{La}/\text{Y}]$ ratio relative to the Galactic halo. The study of the Carina dSph by Venn et al. (2012) also found that the low-metallicity stars have slightly lower $[\text{Y}/\text{Eu}]$ ratios (and slightly higher $[\text{Ba}/\text{Y}]$ ratios) when compared to Galactic stars. The dSph of today are chemically different to the galaxies that merged early during the formation of the Galaxy.

The NS11 study found that the low- α stars show a higher $[\text{Ba}/\text{Y}]$ ratio than the high- α stars. We also find differences for $[\text{Zr}/\text{Fe}]$, $[\text{La}/\text{Zr}]$, $[\text{ls}/\text{Fe}]$, $[\text{hs}/\text{ls}]$, $[\text{Y}/\text{Eu}]$, $[\text{Ba}/\text{Eu}]$, $[\text{ls}/\text{Eu}]$ and $[\text{hs}/\text{Eu}]$ between the two populations. These findings reveal that the low- α population displays many chemical signatures also seen in dSph thereby supporting the accretion hypothesis from NS10.

We now turn our attention to the possible sources that could produce these abundance differences among the neutron-capture elements. Due to the slower rate of chemical evolution, it has been suggested that the dSph and low- α population have been enriched from low-metallicity, low-mass AGB stars (NS11; Venn et al. 2012). Intermediate-mass AGB stars ($>3 M_{\odot}$) produce subsolar $[\text{Ba}/\text{Y}]$ ratios (e.g. Karakas & Lattanzio 2014), i.e. these objects cannot explain the behaviour of $[\text{Ba}/\text{Y}]$ in the low- α population. On the other hand, low-metallicity, low-mass AGB stars ($\lesssim 3 M_{\odot}$) produce $[\text{Ba}/\text{Y}]$ ratios above solar (e.g. Cristallo et al. 2011; Fishlock et al. 2014), and this is consistent with the low- α population. The high $[\text{La}/\text{Zr}]$ and $[\text{hs}/\text{ls}]$ ratios for the low- α population fur-

ther support the hypothesis of enrichment from low-metallicity, low-mass AGBs in the low- α populations as a result of the slower rate of chemical evolution (e.g. in low-mass AGBs, the first s -process peak elements are bypassed in favour of second s -process peak elements).

One issue with this scenario is that we would therefore expect the abundance of Y to be greater than (or equal to) Zr, i.e. $[\text{Y}/\text{Fe}] \geq [\text{Zr}/\text{Fe}]$ (e.g. Bisterzo et al. 2014; Trippella et al. 2016). The observations indicate otherwise and this issue has also been identified in CEMP-s stars (Abate et al. 2015). The reason for the discrepancy between observations and theory is unknown. As noted in Section 3.1, while the abundances of Y and Zr follow each other, there may be a zero-point offset due to errors in the atomic data (e.g. gf values).

If low-metallicity, low-mass stars have contributed to chemical enrichment of the low- α population, then we can place constraints on the enrichment time-scales involved in the earliest dwarf galaxies. Models from Fishlock et al. (2014) predict that low-metallicity, low-mass ($\lesssim 3 M_{\odot}$) stars have lifetimes between 0.29×10^9 and 6.8×10^9 yr.

We also present measurements of Eu; the r -process is responsible for producing the majority of Eu in the Solar system (Snedden et al. 2008). The evolution of $[\text{Eu}/\text{Fe}]$ for both Galactic stars and dSph stars generally follows $[\alpha/\text{Fe}]$ (Venn et al. 2004; Tolstoy et al. 2009) indicating a common origin between α elements and Eu, most likely to be massive stars. If true, then we would expect the low- α stars to have a lower $[\text{Eu}/\text{Fe}]$ abundance than the high- α stars at a given metallicity. However, there is some evidence that the low- α stars have a higher $[\text{Eu}/\text{Fe}]$ ratio than the high- α stars, in contrast to our expectations. This separation is only observed at $[\text{Fe}/\text{H}] < -1$ and includes half the sample. When using the whole metallicity range, the KS test does not find a statistically significant separation between the two populations (see Table 5).

If the low- α stars do have a higher Eu abundance than the high- α stars at lower metallicities, it would suggest that the r -process contributed more to the chemical evolution of the low- α stars at a given metallicity. This is difficult to reconcile with chemical evolution models that support the idea that intense galactic winds remove Eu from a system (Lanfranchi, Matteucci & Cescutti 2008). Another possibility is that the α and Eu abundances are due to different initial mass functions rather than the addition of SNe Ia material (Ishimaru & Wanajo 1999; Travaglio et al. 1999; Kobayashi et al. 2006). To determine if there is the possibility of a statistically significant separation, additional observations of Eu for the NS10 sample are required, particularly for stars with $[\text{Fe}/\text{H}] < -1$.

We observe constant $[\text{Y}/\text{Eu}]$ abundances (~ -0.4 dex) in the lower metallicity low- α stars ($-1.5 \leq [\text{Fe}/\text{H}] \leq -1$) suggesting that Y and Eu have a common origin or time-scale for enrichment, most likely to be the r -process, at these metallicities. The dSph chemical evolution models of Lanfranchi et al. (2008) show constant $[\text{Y}/\text{Eu}]$ ratios up to a $[\text{Fe}/\text{H}]$ ratio of ~ -1.5 . The ratio of $[\text{Y}/\text{Eu}]$ then increases at higher metallicities due to the production of Y by AGB stars.

We observe lower $[\text{Y}/\text{Eu}]$, $[\text{ls}/\text{Eu}]$ and $[\text{hs}/\text{Eu}]$ ratios in the low- α stars compared to the high- α stars with the separation becoming more noticeable at lower metallicities. The study of the Carina dSph by Venn et al. (2012) also found that the low-metallicity stars have slightly lower $[\text{Y}/\text{Eu}]$ ratios (and slightly higher $[\text{Ba}/\text{Y}]$ ratios) when compared to Galactic stars. Venn et al. (2012) concluded that these differences are a result of metal-poor AGB stars contributing to the

⁴ We note that there is some overlap in $[\alpha/\text{Fe}]$ between dSph stars and Galactic halo stars with extreme retrograde orbits (Fulbright 2002; Stephens & Boesgaard 2002). Additionally, the most metal-poor stars in dSph appear to have halo-like $[\alpha/\text{Fe}]$ ratios (Frebel et al. 2010; Norris et al. 2010; Venn et al. 2012).

chemical enrichment of Carina, and this may also apply to the low- α population.

6 CONCLUSIONS

We found that the low- α and high- α populations of NS10 separate when using the abundance ratios of [Sc/Fe], [Zr/Fe], [La/Zr], [ls/Fe], [hs/ls], [Y/Eu], [Ba/Eu], [ls/Eu] and [hs/Eu]. These differences in chemical abundance ratios have only been detected using a differential analysis relative to a reference star with comparable stellar parameters to minimize the errors in abundance measurements. The separation observed in [La/Zr] between the low- and high- α groups confirms the results of NS11 who find a separation in [Ba/Y], where Zr and Y are first *s*-process peak elements and La and Ba are second *s*-process peak elements.

The low abundance of [Y/Eu] found in the low- α population compared to the high- α population matches one of the chemical signatures of present-day dSph. Venn et al. (2004) discard the possibility that the Galactic halo could consist of low-mass dwarf galaxies that have continuously merged with the Galaxy as high [Ba/Y] and low [Y/Eu] ratios are not observed in Galactic halo stars. The idea of more massive mergers with the Galaxy was not ruled out by Venn et al. (2004). In addition, models of AGB stars are unable to explain the lower [hs/Eu] ratios observed in the low- α stars compared to the high- α stars.

The results presented here offer new and important observational evidence regarding the nature of the Galactic halo. It is apparent from the study of NS10 and NS11 that there are at least two populations in the Galactic halo, one that formed *in situ* as well as one accreted from dwarf galaxies. The low [Y/Eu] abundances in dSph have been attributed to the contribution of metal-poor AGB stars (Venn et al. 2004). Low-metallicity AGB stars could also be responsible for the high [Ba/Y] abundances. Only with detailed chemical evolution models, along with measurements of additional *r*-process elements beyond Eu, will there be a better understanding of the chemical enrichment of the earliest dwarf galaxies that merged with the Galactic halo. In particular, through a better understanding of the *s*-process, we will be able to provide constraints on the formation site(s) of the *r*-process.

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